

Accidental Supersymmetry and the Renormalization of Codimension-2 Branes

Matt Williams

August 26, 2013



Based on hep-th/1210.3753, hep-th/1210.5405
with C.P. Burgess, L. van Nierop, S. Parameswaran, and A. Salvio.

6D Chiral, Gauged Supergravity

Salam-Sezgin Solution

SUSY Check (and Preview)

Deforming Salam-Sezgin while Preserving SUSY

Motivation

Brane Tension and the Deficit Angle

Brane Magnetic Charge and Flux Quantization

Application: 4D Vacuum Energy at 1-loop

Classical Result

Example: Scalar Loop

Result: Massive Multiplet

Take-home Message

1. The Salam-Sezgin solution in 6D SUGRA can be deformed by 4D sources in a way that preserves SUSY
2. Surprising result: fields can dynamically adjust to preserve SUSY for *arbitrary 4D source tensions, charges*
3. Such a construction has desirable phenomenological consequences for the naturalness of dark energy

6D Supergravity: Salam-Sezgin Solution

Action for non-trivial bosonic fields:

$$S = - \int d^6 X \sqrt{-g} \left[\frac{1}{2\kappa^2} \left(R + \partial_M \phi \partial^M \phi \right) + \frac{e^{-\phi}}{4g_R^2} F_{MN} F^{MN} + \frac{2g_R^2}{\kappa^4} e^\phi \right]$$

6D Supergravity: Salam-Sezgin Solution

Action for non-trivial bosonic fields:

$$S = - \int d^6 X \sqrt{-g} \left[\frac{1}{2\kappa^2} \left(R + \partial_M \phi \partial^M \phi \right) + \frac{e^{-\phi}}{4g_R^2} F_{MN} F^{MN} + \frac{2g_R^2}{\kappa^4} e^\phi \right]$$

Background:

$$ds^2 = \eta_{\mu\nu} dx^\mu dx^\nu + r^2 (d\theta^2 + \sin^2 \theta d\varphi^2)$$

$$F_{\mu N} = 0, \quad F_{mn} = f \varepsilon_{mn}, \quad \phi = \text{const.}$$

where $r \rightarrow$ radius of the extra-dimensional sphere
and $\varepsilon_{mn} \rightarrow$ extra-dimensional volume form.

6D Supergravity: Salam-Sezgin Solution

Action for non-trivial bosonic fields:

$$S = - \int d^6 X \sqrt{-g} \left[\frac{1}{2\kappa^2} \left(R + \partial_M \phi \partial^M \phi \right) + \frac{e^{-\phi}}{4g_R^2} F_{MN} F^{MN} + \frac{2g_R^2}{\kappa^4} e^\phi \right]$$

Background:

$$ds^2 = \eta_{\mu\nu} dx^\mu dx^\nu + r^2 (d\theta^2 + \sin^2 \theta d\varphi^2)$$
$$F_{\mu N} = 0, \quad F_{mn} = f \varepsilon_{mn}, \quad \phi = \text{const.}$$

where $r \rightarrow$ radius of the extra-dimensional sphere
and $\varepsilon_{mn} \rightarrow$ extra-dimensional volume form.

Salam-Sezgin solution: (Phys. Lett. **B147** (1984) 47)

$$F_{\theta\varphi} = \pm \frac{\varepsilon_{\theta\varphi}}{2r^2} = \pm \frac{\sin \theta}{2}, \quad r^2 e^\phi = \frac{\kappa^2}{4g_R^2}.$$

6D Supergravity: Salam-Sezgin Solution (cont'd)

Supersymmetric?

$$\delta\chi = \frac{1}{\kappa\sqrt{2}}(\partial_M\phi)\Gamma^M\epsilon$$

6D Supergravity: Salam-Sezgin Solution (cont'd)

Supersymmetric?

$$\delta\chi = \frac{1}{\kappa\sqrt{2}}(\partial_M\phi)\Gamma^M\epsilon \quad \checkmark$$

6D Supergravity: Salam-Sezgin Solution (cont'd)

Supersymmetric?

$$\delta\chi = \frac{1}{\kappa\sqrt{2}}(\partial_M\phi)\Gamma^M\epsilon \quad \checkmark$$

$$\delta\lambda = \frac{ie^{-\phi/2}}{2\sqrt{2}g_R} \left(-iF_{MN}\Gamma^{MN} - \frac{4g_R^2}{\kappa^2} e^\phi \right) \epsilon$$

6D Supergravity: Salam-Sezgin Solution (cont'd)

Supersymmetric?

$$\delta\chi = \frac{1}{\kappa\sqrt{2}}(\partial_M\phi)\Gamma^M\epsilon \quad \checkmark$$

$$\delta\lambda = \frac{ie^{-\phi/2}}{2\sqrt{2}g_R} \left(-iF_{MN}\Gamma^{MN} - \frac{4g_R^2}{\kappa^2}e^\phi \right) \epsilon \quad \checkmark$$

$$\rightarrow \epsilon = \begin{pmatrix} \epsilon_{4\pm} \\ 0 \end{pmatrix}, \quad \gamma_5\epsilon_{4\pm} = \pm\epsilon_{4\pm}$$

6D Supergravity: Salam-Sezgin Solution (cont'd)

Supersymmetric?

$$\delta\chi = \frac{1}{\kappa\sqrt{2}}(\partial_M\phi)\Gamma^M\epsilon \quad \checkmark$$

$$\delta\lambda = \frac{ie^{-\phi/2}}{2\sqrt{2}g_R} \left(-iF_{MN}\Gamma^{MN} - \frac{4g_R^2}{\kappa^2} e^\phi \right) \epsilon \quad \checkmark$$

$$\rightarrow \epsilon = \begin{pmatrix} \epsilon_{4\pm} \\ 0 \end{pmatrix}, \quad \gamma_5\epsilon_{4\pm} = \pm\epsilon_{4\pm}$$

$$\delta\psi_M = \frac{\sqrt{2}}{\kappa} D_M\epsilon = \frac{\sqrt{2}}{\kappa} \left(\partial_M - \frac{1}{4}\omega_{MAB}\Gamma^{AB} - iA_M \right) \epsilon$$

6D Supergravity: Salam-Sezgin Solution (cont'd)

Supersymmetric?

$$\delta\chi = \frac{1}{\kappa\sqrt{2}}(\partial_M\phi)\Gamma^M\epsilon \quad \checkmark$$

$$\delta\lambda = \frac{ie^{-\phi/2}}{2\sqrt{2}g_R} \left(-iF_{MN}\Gamma^{MN} - \frac{4g_R^2}{\kappa^2} e^\phi \right) \epsilon \quad \checkmark$$

$$\rightarrow \epsilon = \begin{pmatrix} \epsilon_{4\pm} \\ 0 \end{pmatrix}, \quad \gamma_5\epsilon_{4\pm} = \pm\epsilon_{4\pm}$$

$$\delta\psi_M = \frac{\sqrt{2}}{\kappa} D_M\epsilon = \frac{\sqrt{2}}{\kappa} \left(\partial_M - \frac{1}{4}\omega_{MAB}\Gamma^{AB} - iA_M \right) \epsilon \quad \checkmark$$

$$\rightarrow \partial_M\epsilon = 0$$

6D Supergravity: Salam-Sezgin Solution (cont'd)

Supersymmetric?

$$\delta\chi = \frac{1}{\kappa\sqrt{2}}(\partial_M\phi)\Gamma^M\epsilon \quad \checkmark$$

$$\delta\lambda = \frac{ie^{-\phi/2}}{2\sqrt{2}g_R} \left(-iF_{MN}\Gamma^{MN} - \frac{4g_R^2}{\kappa^2}e^\phi \right) \epsilon \quad \checkmark$$

$$\rightarrow \epsilon = \begin{pmatrix} \epsilon_{4\pm} \\ 0 \end{pmatrix}, \quad \gamma_5\epsilon_{4\pm} = \pm\epsilon_{4\pm}$$

$$\delta\psi_M = \frac{\sqrt{2}}{\kappa}D_M\epsilon = \frac{\sqrt{2}}{\kappa} \left(\partial_M - \frac{1}{4}\omega_{MAB}\Gamma^{AB} - iA_M \right) \epsilon \quad \checkmark$$

$$\rightarrow \partial_M\epsilon = 0$$

$$\left(\omega_{\varphi 45} = \cos\theta - b, \quad A_\varphi = \mp\frac{1}{2}(\cos\theta - b) \right)$$

$$(b = \pm 1 \leftarrow \text{N/S pole})$$

6D Supergravity: Salam-Sezgin Solution (cont'd)

Supersymmetric?

$$\delta\chi = \frac{1}{\kappa\sqrt{2}}(\partial_M\phi)\Gamma^M\epsilon \quad \checkmark$$

$$\delta\lambda = \frac{ie^{-\phi/2}}{2\sqrt{2}g_R} \left(-iF_{MN}\Gamma^{MN} - \frac{4g_R^2}{\kappa^2}e^\phi \right) \epsilon \quad \checkmark$$

$$\rightarrow \epsilon = \begin{pmatrix} \epsilon_{4\pm} \\ 0 \end{pmatrix}, \quad \gamma_5\epsilon_{4\pm} = \pm\epsilon_{4\pm}$$

$$\delta\psi_M = \frac{\sqrt{2}}{\kappa}D_M\epsilon = \frac{\sqrt{2}}{\kappa} \left(\partial_M - \frac{1}{4}\omega_{MAB}\Gamma^{AB} - iA_M \right) \epsilon \quad \checkmark$$

$$\rightarrow \partial_M\epsilon = 0$$

$$\left(\omega_{\varphi 45} = \alpha \cos\theta - b, \quad A_\varphi = \mp \frac{\alpha}{2}(\cos\theta - b) + b\Phi_b \right)$$

$$\rightarrow \Phi_b = \pm \frac{1}{2}(1 - \alpha)$$

6D Supergravity: Salam-Sezgin Solution (cont'd)

Supersymmetric?

$$\delta\chi = \frac{1}{\kappa\sqrt{2}}(\partial_M\phi)\Gamma^M\epsilon \quad \checkmark$$

$$\delta\lambda = \frac{ie^{-\phi/2}}{2\sqrt{2}g_R} \left(-iF_{MN}\Gamma^{MN} - \frac{4g_R^2}{\kappa^2}e^\phi \right) \epsilon \quad \checkmark$$

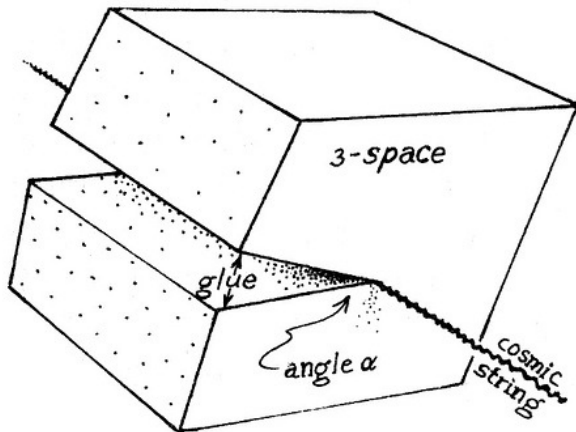
$$\rightarrow \epsilon = \begin{pmatrix} \epsilon_{4\pm} \\ 0 \end{pmatrix}, \quad \gamma_5\epsilon_{4\pm} = \pm\epsilon_{4\pm}$$

$$\delta\psi_M = \frac{\sqrt{2}}{\kappa}D_M\epsilon = \frac{\sqrt{2}}{\kappa} \left(\partial_M - \frac{1}{4}\omega_{MAB}\Gamma^{AB} - iA_M \right) \epsilon \quad \times$$

$$\rightarrow \partial_M\epsilon = 0$$

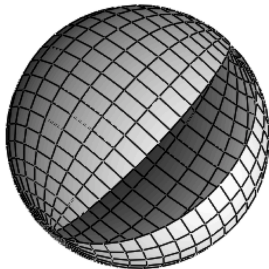
$$\left(\omega_{\varphi 45} = \alpha \cos\theta - b, \quad A_\varphi = \mp \frac{\alpha}{2}(\cos\theta - b) \right)$$

Motivation: Analogy with Cosmic Strings



Picture credit: R. Penrose, *The Road to Reality*

Rugby-Ball Background



$$S_b \supset - \int d^4x \sqrt{-g_4} T_b, \quad b = \pm 1.$$

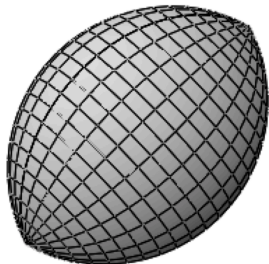
Background Metric:

$$ds^2 = g_{\mu\nu} dx^\mu dx^\nu + r^2 (d\theta^2 + \alpha^2 \sin^2 \theta d\varphi^2)$$

where

$$\alpha = 1 - \frac{\kappa^2 T_+}{2\pi} = 1 - \frac{\kappa^2 T_-}{2\pi}.$$

Deficit angle: $\delta = 2\pi(1 - \alpha)$.



Brane Magnetic Charge

$$S_b \supset \frac{\mathcal{A}_b}{2g_R^2} \int *F = \int d^4x \sqrt{-g_4} \frac{\mathcal{A}_b}{2g_R^2} \varepsilon^{mn} F_{mn} .$$
$$\left(\mathcal{L}_{gf} = -\frac{e^{-\phi}}{4g_R^2} F_{MN} F^{MN} \right)$$

Alters gauge field boundary conditions ($b = \pm 1 \leftarrow$ N/S pole):

$$A_\varphi \Big|_{\cos \theta = b} = b \frac{e^\phi \mathcal{A}_b}{2\pi} := b \Phi_b .$$

Flux Quantization

Non-local condition:

$$N = \sum_b \Phi_b + \frac{1}{2\pi} \int d^2y F_{\theta\varphi}.$$

Gauge field strength:

$$F_{\theta\varphi} = \pm \frac{\varepsilon_{\theta\varphi}}{2r^2} = \pm \frac{\alpha \sin \theta}{2};$$

$$\rightarrow \sum_b \Phi_b = \pm(1 - \alpha).$$

Identical branes:

$$\Phi_b = \pm \frac{1}{2}(1 - \alpha) \quad \text{or} \quad e^{\phi_b} \mathcal{A}_b = \pm \frac{\kappa^2 T_b}{2}.$$

Application: 4D Vacuum Energy at 1-loop

Taking $\mathcal{A}_+ \neq \mathcal{A}_-$ provides a framework to check the technical naturalness of the classical result:

$$\rho_V \Big|_{\text{class.}} = \frac{1}{2} \frac{\partial \mathcal{L}_b}{\partial \phi} = 0.$$

(L. van Nierop and C.P. Burgess in hep-th/1108.0345.)

1-loop Vacuum Energy Example — Scalar Loop

1-loop Vacuum Energy Example — Scalar Loop

Action: $S_B \supset - \int d^6 X \sqrt{-g} \left(\frac{1}{2} g^{MN} D_M \psi D_N \psi + \frac{1}{2} m^2 \psi^2 \right)$ where

$$D_M \psi = \partial_M \psi - iq A_M \psi.$$

1-loop Vacuum Energy Example — Scalar Loop

Action: $S_B \supset - \int d^6 X \sqrt{-g} \left(\frac{1}{2} g^{MN} D_M \psi D_N \psi + \frac{1}{2} m^2 \psi^2 \right)$ where

$$D_M \psi = \partial_M \psi - iq A_M \psi.$$

Since $\Sigma_{1L} = - \int d^4 x V_{1L} = \frac{i}{2} \ln \text{Det}(-D^M D_M + m^2)$,

$$\rightarrow V_{1L} = \frac{1}{2} \mu^{4-d} \sum_{jn} \int \frac{d^d k_E}{(2\pi)^d} \ln \left(\frac{k_E^2 + m_{jn}^2 + m^2}{\mu^2} \right).$$

1-loop Vacuum Energy Example — Scalar Loop

Action: $S_B \supset - \int d^d X \sqrt{-g} \left(\frac{1}{2} g^{MN} D_M \psi D_N \psi + \frac{1}{2} m^2 \psi^2 \right)$ where

$$D_M \psi = \partial_M \psi - iq A_M \psi.$$

Since $\Sigma_{1L} = - \int d^d x V_{1L} = \frac{i}{2} \ln \text{Det}(-D^M D_M + m^2)$,

$$\rightarrow V_{1L} = \frac{1}{2} \mu^{4-d} \sum_{jn} \int \frac{d^d k_E}{(2\pi)^d} \ln \left(\frac{k_E^2 + m_{jn}^2 + m^2}{\mu^2} \right).$$

Heat kernel expansion:

$$V_{1L} = - \frac{\mu^{4-d}}{2(4\pi r^2)^{d/2}} \int_0^\infty \frac{dt}{t^{1+d/2}} e^{-t(mr)^2} \left(\sum_{j,n} e^{-t(m_{jn}r)^2} \right).$$

Technical Details

- ▶ Obtain KK spectrum for a scalar field (or spin- $\frac{1}{2}$, etc.)
- ▶ Perform the sum over KK modes
- ▶ Extract the UV divergences by taking $\lim_{d \rightarrow 4} V_{1L}$

Technical Details

- ▶ Obtain KK spectrum for a scalar field (or spin- $\frac{1}{2}$, etc.)
- ▶ Perform the sum over KK modes
- ▶ Extract the UV divergences by taking $\lim_{d \rightarrow 4} V_{1L}$

Final result:

$$V_{1L} = \frac{\mathcal{C}}{(4\pi)^2} \left[\frac{1}{4-d} + \ln\left(\frac{\mu}{m}\right) \right] + \dots$$

where

$$\mathcal{C} = \frac{s_{-1}}{6} m^6 r^2 - \frac{s_0}{2} m^4 + s_1 \frac{m^2}{r^2} - \frac{s_2}{r^4}.$$

Property:

$$s_i = \alpha s_i^{\text{sph}} + \sum_b \delta s_{i(b)}(\alpha, \Phi_b).$$

Technical Details

- ▶ Obtain KK spectrum for a scalar field (or spin- $\frac{1}{2}$, etc.)
- ▶ Perform the sum over KK modes
- ▶ Extract the UV divergences by taking $\lim_{d \rightarrow 4} V_{1L}$

Final result:

$$V_{1L} = \frac{\mathcal{C}}{(4\pi)^2} \left[\frac{1}{4-d} + \ln\left(\frac{\mu}{m}\right) \right] + \dots$$

where

$$\mathcal{C} = \frac{s_{-1}}{6} m^6 r^2 - \frac{s_0}{2} m^4 + s_1 \frac{m^2}{r^2} - \frac{s_2}{r^4}.$$

Property:

$$s_i = \alpha s_i^{\text{sph}} + \sum_b \delta s_{i(b)}(\alpha, \Phi_b).$$

- ▶ Identify the requisite bulk, brane counterterms
- ▶ Obtain beta functions for renormalized quantities

Result: Massive Multiplet

$$\text{Back-reaction: } \rho_{V(1L)} = -\frac{1}{2} \frac{\partial V_{R(1L)}}{\partial \phi} ;$$

- ▶ Massive multiplet: ($\eta = |\Delta\Phi|/(2\alpha)$, $m^2(\phi) = M^2 e^\phi$)

$$\rho_{V(1L)} = \frac{\eta}{(4\pi r^2)^2} \left[\left(\frac{\kappa M}{2g_R} \right)^4 - \left(\frac{1}{3} + \frac{(1-\alpha^2)}{2\alpha} \eta - \frac{\eta^2}{3} \right) \left(\frac{\kappa M}{2g_R} \right)^2 \right] \ln \left(\frac{M_s}{M} \right) .$$

Result: Massive Multiplet

Back-reaction: $\rho_{V(1L)} = -\frac{1}{2} \frac{\partial V_{R(1L)}}{\partial \phi}$;

- ▶ Massive multiplet: ($\eta = |\Delta\Phi|/(2\alpha)$, $m^2(\phi) = M^2 e^\phi$)

$$\rho_{V(1L)} = \frac{\eta}{(4\pi r^2)^2} \left[\left(\frac{\kappa M}{2g_R} \right)^4 - \left(\frac{1}{3} + \frac{(1-\alpha^2)\eta - \eta^2}{2\alpha} \right) \left(\frac{\kappa M}{2g_R} \right)^2 \right] \ln \left(\frac{M_s}{M} \right) .$$

If brane charges are balanced (*i.e.* $\eta = 0$), then $\rho_{V(1L)} = 0$.

Result: Massive Multiplet

Back-reaction: $\rho_{V(1L)} = -\frac{1}{2} \frac{\partial V_{R(1L)}}{\partial \phi}$;

- ▶ Massive multiplet: ($\eta = |\Delta\Phi|/(2\alpha)$, $m^2(\phi) = M^2 e^\phi$)

$$\rho_{V(1L)} = \frac{\eta}{(4\pi r^2)^2} \left[\left(\frac{\kappa M}{2g_R} \right)^4 - \left(\frac{1}{3} + \frac{(1-\alpha^2)}{2\alpha} \eta - \frac{\eta^2}{3} \right) \left(\frac{\kappa M}{2g_R} \right)^2 \right] \ln \left(\frac{M_s}{M} \right) .$$

If brane charges are balanced (*i.e.* $\eta = 0$), then $\rho_{V(1L)} = 0$.

Technically natural? ✓

Take home message:

1. The Salam-Sezgin solution in 6D SUGRA can be deformed by 4D sources in a way that preserves SUSY
2. Surprising result: fields can dynamically adjust to preserve SUSY for *arbitrary 4D source tensions, charges*
3. Such a construction has desirable phenomenological consequences for the naturalness of dark energy

Take home message:

1. The Salam-Sezgin solution in 6D SUGRA can be deformed by 4D sources in a way that preserves SUSY
2. Surprising result: fields can dynamically adjust to preserve SUSY for *arbitrary 4D source tensions, charges*
3. Such a construction has desirable phenomenological consequences for the naturalness of dark energy

Thank you for your attention!

For more information:

- ▶ M. Williams, C. P. Burgess, L. van Nierop and A. Salvio, "Running with Rugby Balls: Bulk Renormalization of Codimension-2 Branes," JHEP **1301**, 102 (2013) [arXiv:1210.3753 [hep-th]]
- ▶ C. P. Burgess, L. van Nierop, S. Parameswaran, A. Salvio and M. Williams, "Accidental SUSY: Enhanced Bulk Supersymmetry from Brane Back-reaction," JHEP **1302**, 120 (2013) [arXiv:1210.5405 [hep-th]]

Example — Scalar Loop Calc. (cont'd)

- ▶ Solve for scalar's KK spectrum, $m_{jn}^2(\alpha, N, \Phi_b)$:

$$m_{jn}^2 = \frac{1}{r^2} \left[(j + y + z)(j + y + z + 1) - \frac{\mathcal{N}^2}{4} \right]$$

where $y = \frac{1}{2\alpha}|n - \Phi_+|$, $z = \frac{1}{2\alpha}|n - N + \Phi_-|$, and $\mathcal{N} = \frac{N - \Phi}{\alpha}$.

(When $1 - \alpha = N = \Phi_b = 0$, $m_{jn}^2 = \frac{\ell(\ell+1)}{r^2}$ with $\ell := j + |n|$.)

- ▶ Sum over KK spectrum using E.-M., Poisson resummation:

$$\sum_{j,n} e^{-t(m_{jn}r)^2} = \frac{s_{-1}}{t} + s_0 + s_1 t + s_2 t^2 + \mathcal{O}(t^2);$$

- ▶ Take limit $d \rightarrow 4$:

$$V_{1L} = \frac{\mathcal{C}}{(4\pi)^2} \left[\frac{1}{4-d} + \ln\left(\frac{\mu}{m}\right) \right] + \mathcal{V}_f,$$

where

$$\mathcal{C} := \frac{s_{-1}}{6} m^6 r^2 - \frac{s_0}{2} m^4 + s_1 \frac{m^2}{r^2} - \frac{s_2}{r^4}.$$

“Bulk Renorm. is Independent of BC’s”

Sphere, no BLF's ($\alpha = 1, \Phi_b = 0$): (Kantowski & Milton, 1987)

$$s_{-1}^{\text{sph}} = 1, \quad s_0^{\text{sph}} = \frac{1}{3}, \quad s_1^{\text{sph}} = \frac{1}{15} - \frac{N^2}{24}, \quad s_2^{\text{sph}} = \frac{4}{315} - \frac{N^2}{40}.$$

“Bulk Renorm. is Independent of BC’s”

Sphere, no BLF's ($\alpha = 1, \Phi_b = 0$): (Kantowski & Milton, 1987)

$$s_{-1}^{\text{sph}} = 1, \quad s_0^{\text{sph}} = \frac{1}{3}, \quad s_1^{\text{sph}} = \frac{1}{15} - \frac{N^2}{24}, \quad s_2^{\text{sph}} = \frac{4}{315} - \frac{N^2}{40}.$$

Required bulk counterterms:

$$\mathcal{L}_{1Lct} = - \left[U + \frac{1}{2\kappa^2} R + \frac{\zeta_{R^2}}{\kappa} \bar{R}^2 + \frac{1}{4g^2} F_{MN} F^{MN} + \zeta_{R^3} \bar{R}^3 + \frac{\kappa \zeta_{AR}}{8g^2} R F_{MN} F^{MN} \right]$$

so that $V_{ct} = - \int d^2y \mathcal{L}_{ct} \Big|_{\text{bkgd}}$ cancels all UV divergences.

“Bulk Renorm. is Independent of BC’s”

Sphere, no BLF’s ($\alpha = 1, \Phi_b = 0$): (Kantowski & Milton, 1987)

$$s_{-1}^{\text{sph}} = 1, \quad s_0^{\text{sph}} = \frac{1}{3}, \quad s_1^{\text{sph}} = \frac{1}{15} - \frac{N^2}{24}, \quad s_2^{\text{sph}} = \frac{4}{315} - \frac{N^2}{40}.$$

Required bulk counterterms:

$$\mathcal{L}_{1\text{Lct}} = - \left[U + \frac{1}{2\kappa^2} R + \frac{\zeta_{R^2}}{\kappa} \bar{R}^2 + \frac{1}{4g^2} F_{MN} F^{MN} + \zeta_{R^3} \bar{R}^3 + \frac{\kappa \zeta_{AR}}{8g^2} R F_{MN} F^{MN} \right]$$

so that $V_{\text{ct}} = - \int d^2y \mathcal{L}_{\text{ct}} \Big|_{\text{bkgd}}$ cancels all UV divergences.

Running:

$$\begin{aligned} \mu \frac{\partial U}{\partial \mu} &= -\frac{m^6}{6(4\pi)^3}, & \mu \frac{\partial}{\partial \mu} \left(\frac{1}{\kappa^2} \right) &= -\frac{m^4}{6(4\pi)^3}, & \mu \frac{\partial}{\partial \mu} \left(\frac{\zeta_{R^2}}{\kappa} \right) &= -\frac{m^2}{60(4\pi)^3}, \\ \mu \frac{\partial}{\partial \mu} \left(\frac{1}{g^2} \right) &= \frac{2q^2 m^2}{3(4\pi)^3}, & \mu \frac{\partial \zeta_{R^3}}{\partial \mu} &= -\frac{1}{630(4\pi)^3}, & \mu \frac{\partial}{\partial \mu} \left(\frac{\kappa \zeta_{AR}}{g^2} \right) &= \frac{2q^2}{5(4\pi)^3}. \end{aligned}$$

The General Case

$$s_{-1}^s = \alpha,$$

$$s_0^s(\omega, N, \Phi_b) = \alpha \left[\frac{1}{6} + \frac{\omega^2}{6} (1 - 3F) \right],$$

$$s_1^s(\omega, N, \Phi_b) = \alpha \left[\frac{1}{180} - \frac{\mathcal{N}^2}{24} + \frac{\omega^2}{18} (1 - 3F) - \frac{\omega^3 \mathcal{N}}{12} \sum_b \Phi_b G(|\Phi_b|) + \frac{\omega^4}{180} (1 - 15F^{(2)}) \right],$$

$$s_2^s(\omega, N, \Phi_b) = \alpha \left[-\frac{1}{504} - \frac{11\mathcal{N}^2}{720} + \left(\frac{1}{90} - \frac{\mathcal{N}^2}{144} \right) (1 - 3F)\omega^2 - \frac{\omega^3 \mathcal{N}}{24} \sum_b \Phi_b G(|\Phi_b|) \right. \\ \left. + \frac{\omega^4 (1 - \mathcal{N}^2)}{360} (1 - 15F^{(2)}) - \frac{\omega^5 \mathcal{N}}{120} \sum_b \Phi_b G(|\Phi_b|) (1 + 3F_b) \right. \\ \left. + \left(\frac{1}{1260} - \frac{F^{(2)}}{120} - \frac{F^{(3)}}{60} \right) \omega^6 \right].$$

with

$$\omega := \alpha^{-1}, \quad \mathcal{N} := \frac{N - \Phi}{\alpha}, \quad F_b := |\Phi_b| (1 - |\Phi_b|), \quad F^{(n)} := \sum_b F_b^n,$$

$$F := F^{(1)}, \quad G(|\Phi_b|) := (1 - |\Phi_b|)(1 - 2|\Phi_b|).$$

“Brane Renorm. Depends Only on Itself”

- ▶ Bulk renormalization is independent of brane physics — contribution to V_{1L} scales with α (volume factor);
- ▶ Brane renormalization vanishes in the limit $\alpha \rightarrow 1$, $\Phi_b \rightarrow 0$.

$$s_i = \alpha s_i^{\text{sph}} + \sum_b \delta s_{i(b)}$$

“Brane Renorm. Depends Only on Itself”

- ▶ Bulk renormalization is independent of brane physics — contribution to V_{1L} scales with α (volume factor);
- ▶ Brane renormalization vanishes in the limit $\alpha \rightarrow 1$, $\Phi_b \rightarrow 0$.

$$s_i = \alpha s_i^{\text{sph}} + \sum_b \delta s_{i(b)}$$

Brane counterterms:

$$\mathcal{L}_{b,\text{ct}} = -\sqrt{-g_4} \left(T_{b,\text{ct}} - \frac{\mathcal{A}_{b,\text{ct}}}{2g_R^2} \epsilon \cdot F + \frac{\zeta_{Rb}}{\kappa} R + \frac{\kappa \zeta_{Ab}}{4g_R^2} F^2 + \frac{\kappa \zeta_{\tilde{A}Rb}}{2g_R^2} R \epsilon \cdot F + \zeta_{R^2b} \overline{R^2} \right).$$

“Brane Renorm. Depends Only on Itself”

- ▶ Bulk renormalization is independent of brane physics — contribution to V_{1L} scales with α (volume factor);
- ▶ Brane renormalization vanishes in the limit $\alpha \rightarrow 1$, $\Phi_b \rightarrow 0$.

$$s_i = \alpha s_i^{\text{sph}} + \sum_b \delta s_{i(b)}$$

Brane counterterms:

$$\mathcal{L}_{b,\text{ct}} = -\sqrt{-g_4} \left(T_{b,\text{ct}} - \frac{\mathcal{A}_{b,\text{ct}}}{2g_R^2} \epsilon \cdot F + \frac{\zeta_{Rb}}{\kappa} R + \frac{\kappa \zeta_{Ab}}{4g_R^2} F^2 + \frac{\kappa \zeta_{\tilde{A}Rb}}{2g_R^2} R \epsilon \cdot F + \zeta_{R^2b} \overline{R}^2 \right).$$

Example: ($\omega := \alpha^{-1}$, $\delta\omega := \omega - 1$)

$$\begin{aligned} \mu \frac{\partial T_b}{\partial \mu} &= \frac{m^4}{2(4\pi)^2 \omega} \left(\frac{\delta\omega}{6} + \frac{\delta\omega^2}{12} - \frac{\omega^2}{2} |\Phi_b| (1 - |\Phi_b|) \right), \\ \mu \frac{\partial}{\partial \mu} \left(\frac{\mathcal{A}_b}{g_R^2} \right) &= -\frac{\Phi_b \omega^2 m^2}{6(4\pi)^2} (1 - |\Phi_b|) (1 - 2|\Phi_b|). \end{aligned}$$